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LETTER TO THE EDITOR

GAS in Protoplanetary Systems (GASPS)^{*}

I. First results

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ABSTRACT

Context. Circumstellar discs are ubiquitous around young stars, but rapidly dissipate their gas and dust on timescales of a few Myr. The *Herschel* Space Observatory allows for the study of the warm disc atmosphere, using far-infrared spectroscopy to measure gas content and excitation conditions, and far-IR photometry to constrain the dust distribution.

Aims. We aim to detect and characterize the gas content of circumstellar discs in four targets as part of the *Herschel* science demonstration phase.

Methods. We carried out sensitive medium resolution spectroscopy and high sensitivity photometry at $\lambda \sim 60\text{--}190\ \mu\text{m}$ using the Photodetector Array Camera and Spectrometer instrument on the *Herschel* Space Observatory.

Results. We detect [OI] 63 μm emission from the young stars HD 169142, TW Hydrae, and RECX 15, but not HD 181327. No other lines, including [CII] 158 and [OI] 145, are significantly detected. All four stars are detected in photometry at 70 and 160 μm . Extensive models are presented in associated papers.

Key words. stars: pre-main sequence – protoplanetary disks – infrared: stars

1. Introduction

Circumstellar discs of gas and dust are the environments in which planets form; thus, studies of the evolution of these discs will inform our understanding of the formation of gas giants. The *Herschel* Space Observatory (Pilbratt et al. 2010), with its facilities for sensitive spectrometry and photometry, allows for significantly deeper searches for circumstellar gas than were previously possible, providing tighter constraints on the timescale of giant planet formation and the mechanisms of circumstellar disc evolution.

The GAS in Protoplanetary Systems *Herschel* open time key project (GASPS, P.I. W.R.F. Dent, see Dent et al., in prep.) aims to characterize the evolution of gas and dust in circumstellar discs. Using the Photodetector Array Camera and Spectrometer (PACS, Poglitsch et al. 2010) integral field unit spectrometer and bolometer array, we are carrying out a 400 hour survey of 240 young stars for gas line emission and sensitive far-IR photometry. We focus on three primary emission lines, [OI] 63.184 μm , [OI] 145.525 μm , and [CII] 157.741 μm , as well as several emission lines of H₂O and CO and broadband photometry at 70 and

160 μm to measure dust emission. Our survey covers young stars at ages from 1–30 Myr, in the Taurus, Upper Scorpius, η Chamaeleontis, TWA, Beta Pictoris, and Tucana Horologium associations, as well as HAeBe stars at a variety of ages 0.1–30 Myr, and with masses ranging from $\sim 0.5\text{--}4\ M_{\odot}$. We chose systems known from near-IR and millimeter photometry and accretion indicators to have discs in a variety of states, from those still accreting to debris discs, as well as a sample of stars showing no signs of a disc. We focus on optically-visible stars with low extinction and small envelope masses, i.e. SED Class II–III.

At a basic level, observations of [CII] and [OI] will provide a determination of the presence of gas. Comparison to photodissociation models (e.g. Kaufman et al. 1999) allows for the use of [CII] and [OI] emission to place constraints on the density and temperature structure of circumstellar discs. Such observations will help illuminate disparities between disc mass estimates based on sub-mm continuum measurements of dust emission and gas tracers such as CO. However, atomic C and O have emission lines in the far-IR that are largely unobservable from the ground. While ISO allowed for studies of atomic emission from a few bright sources (e.g. Lorenzetti et al. 1999), the *Herschel* PACS instrument has 2 orders of magnitude better point-source sensitivity than the ISO *Long Wavelength Spectrograph*, opening up the study of circumstellar disc gas to hundreds more sources.

^{*} *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

To aid with the interpretation of our PACS line and continuum data, we have interfaced tools for calculating dust and line radiative transfer (MCFOST, Pinte et al. 2006) and gas thermo-chemistry (ProDiMo, Woitke et al. 2009). We have used these codes to produce a grid of 300 000 disc models, the Disc Evolution with Neat Theory (DENT, Woitke et al. 2010).

In this paper, we present the first results of science demonstration phase observations of a Herbig Ae star (HD 169142), two T-Tauri stars (RECX 15 in η Cha and TW Hya), and a debris disc system (HD 181327). All have ages in the range 6–12 Myr, and bear discs that have undergone some evolution. Near/mid-IR photometry and mm photometry have indicated disc clearing or grain growth in the inner discs of HD 169142 and TW Hya (e.g. Grady et al. 2007; Hughes et al. 2007). Mass accretion rate estimates range from $<7 \times 10^{-10}$ for HD169142 (Grady et al. 2007) to 5×10^{-10} – $2 \times 10^{-9} M_{\odot}/\text{yr}$ for TW Hya (Muzerolle et al. 2000; Herczeg et al. 2004); and $\sim 10^{-9} M_{\odot}/\text{yr}$ for RECX 15 (Lawson et al. 2002).

In Sect. 2, we describe our observations and data reduction, and present our results. We present a preliminary discussion of the implications for the states of these discs in Sect. 3.

2. Observations and results

For our SDP targets, we carried out both spectroscopy and photometry with the PACS instrument on *Herschel*. For each target, we carried out a 1669 s PacsLineSpec observation, a 5150 s PacsRangeSpec observation, and either a 133 s PacsPhoto scan map (RECX 15 and TW Hya) or 159 s PacsPhoto point source observation (HD 169142 and HD 181327). We carried out an additional 220 s scan map observation of RECX 15, to test observation modes. We carried out the reduction of our observations with the facility reduction tool, HIPE (Ott 2010), using the scripts disseminated at the *Herschel* data reduction workshop in January, 2010¹.

2.1. Spectroscopy

Spectroscopic observations with the PACS IFU were carried out in chop-nod mode in order to remove the telescope emission. The PACS IFU is a 5×5 filled array, with $9.4'' \times 9.4''$ spaxels. Observations were carried out such that the central IFU spaxel was centered on the target in both nods. Our SDP observations were carried out with a $2''$ dither. Each PacsLineSpec observation includes simultaneous observations from 62.93–64.43 and 180.76–190.29 μm . Our PacsRangeSpec observations included, in series, simultaneous observations from 71.81–73.28 and 143.59–146.53 μm , 78.37–79.76 and 156.70–159.47 μm , and 89.28–90.48 and 178.51–180.96 μm . Typical continuum rms for all regions are ~ 0.3 Jy, with 1σ line flux uncertainties of 2 – $5 \times 10^{-18} \text{ W/m}^2$. Spectral resolution ranges from $R \sim 1100$ –3400. Full details of the observations are given in Dent et al. (in prep).

For our point source flux measurements, we extracted the A and B nods for the central IFU spaxel. For each nod, we then Nyquist binned in wavelength with non-overlapping bins half the width of the instrumental resolution. The mean of the two nods restores the flux of the objects. We apply wavelength dependent flux and aperture corrections released by the PACS development team², and assume a flux calibration uncertainty of 40%.

¹ <https://nhscsci.ipac.caltech.edu/sc/index.php/Workshops/>

² PacsSpectroscopyPerformanceAndCalibration_v1_1_0.pdf

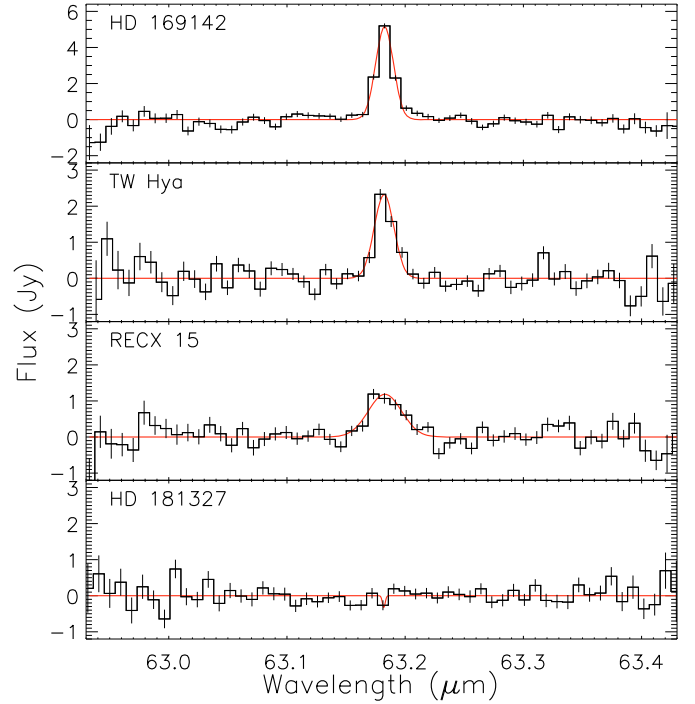


Fig. 1. Baseline subtracted, aperture and flux corrected [OI] 63 μm spectra. Error bars indicate 1σ statistical errors; gaussian line fits are shown in red.

To extract line fluxes and measure continuum flux from our reduced, nod-combined, flux corrected spectra, we carried out an inverse-error weighted fit of a gaussian and first-order polynomial to the spectra. Using the polynomial fit to each spectral region, we estimated the continuum emission at the rest wavelength for each line. To estimate the error on the continuum, we calculated the error of the mean for the residual of the polynomial fit in a region from 2 to 10 instrumental $FWHM$ from the line rest wavelength. For emission lines, we report the flux of detected lines as the integrated flux of the gaussian line fit. We calculate 1σ fluxes as the integral of a gaussian with height equal to the continuum RMS, and width equal to the instrumental $FWHM$; we report 3σ upper limits for non-detections.

The [OI] 63 μm emission line is clearly detected towards HD 169142, RECX 15, and TW Hya from our SDP sample; with $FWHM \sim 90 \text{ km s}^{-1}$, HD 169142 and TW Hya are spectrally unresolved. The emission from RECX 15, on the other hand, appears much broader, with a $FWHM \sim 150 \text{ km s}^{-1}$; this is a marginal result and could be due to off-center or low S/N observations in the two nods (further discussion in Woitke et al., in prep.). We show our baseline subtracted [OI] 63 μm spectra in Fig. 1. Examination of spectra from the off-center data cubes show no indications of extended or field emission. In Table 1, we show the [OI] 63, [OI] 145, and [CII] 157 fluxes and 3σ upper limits, as well as the measured continuum fluxes and 3σ upper limits. At this stage, no other lines were detected with high significance.

2.2. Photometry

Simultaneous 70 and 160 μm continuum observations were obtained with the PACS photometer. The blue camera consists of a 64×32 array of $3.2''$ pixels, while the red camera consists of a 32×16 array of $6.4''$ pixels. Point source mode consists of 50'' chod-nod observations for sky subtraction and a 3 point dither ($8.5''$) for higher spatial sampling. Scan maps were carried out with a $20''/\text{s}$ scan rate, with a 10 Hz integration cycle.

Table 1. Stellar properties and observational results.

Object	SpT	Age Myr	Distance pc	[OI] 63 ^a 10 ⁻¹⁸ W/m ²	cont. ^{a,b} Jy	[OI] 145 ^a 10 ⁻¹⁸ W/m ²	cont. ^{a,c} Jy	[CII] 157 ^a 10 ⁻¹⁸ W/m ²	cont. ^{a,d} Jy	70 μ m ^e Jy	160 μ m ^f Jy
HD 169142	A5 ⁽¹⁾	6 ⁺⁶⁽²⁾ ₋₃	145 ⁽³⁾	71.7 \pm 3.8	18.89	<10.4	13.76	<6.4	14.49	27.35 \pm 0.03	17.39 \pm 0.05
HD 181327	F5 ⁽⁴⁾	12 ⁺⁸⁽⁵⁾ ₋₄	50.6 ⁽⁶⁾	<9.3	1.75	<8.0	0.80	<8.5	0.77	1.90 \pm 0.03	0.87 \pm 0.03
TW Hya	K7e ⁽⁷⁾	10 ⁺¹⁰⁽⁸⁾ ₋₇	56 ⁽⁹⁾	36.5 \pm 3.6	2.99	<5.5	7.00	<6.6	8.79	3.90 \pm 0.02	7.38 \pm 0.04
RECX 15	M2e ⁽¹⁰⁾	9 ⁺²⁽¹¹⁾ ₋₂	97 ⁽¹²⁾	31.5 \pm 3.1	0.14	<4.6	0.10	<7.4	<0.09	0.18 \pm 0.02 ^g	0.07 \pm 0.01 ^g

Notes. (a) Calibration uncertainty 40%; (b) $1\sigma \sim 0.03$ – 0.05 Jy; (c) $1\sigma \sim 0.02$ – 0.04 Jy; (d) $1\sigma \sim 0.03$ – 0.05 Jy; (e) Calibration uncertainty 5%; (f) Calibration uncertainty 10%; (g) Results from two stacked scan maps.

References. (1) Dunkin et al. (1997); (2) Grady et al. (2007); (3) van Boekel et al. (2005); (4) Nordstrom et al. (2004); (5) Zuckerman & Song (2004); (6) Perryman et al. (1997); (7) de la Reza et al. (1989); (8) Barrado Y Navascués (2006); (9) Wichmann et al. (1998); (10, 11) Lawson et al. (2002); (12) Mamajek et al. (2000)

All sources appear spatially unresolved, with $FWHM \sim 5''$ and $11''$ at 70 and 160 μ m, respectively. We carried out photometry on images at the instrumental pixel size using a $21''$ aperture for both blue and red camera observations. The PACS team has reported flux corrections and aperture corrections³, resulting in a flux calibration uncertainty of 5% in the blue camera, and 10% in the red camera. All four of our SDP targets are detected in both bands; we report these continuum measurements without color correction in Table 1. None of the sources show signs of extended emission. Comparison with publicly available IRAS and *Spitzer* photometry shows a general agreement within the systematic uncertainties. Our model fitting efforts of lines and continuum from HD 169142 and TW Hya, using MCFOST and ProDiMo, are presented in Meeus et al. (2010) and Thi et al. (2010); RECX 15 and HD 181327 are addressed in further detail in Woitke et al. (in prep) and Ménard et al. (in prep)

We find one exception to a good agreement between continuum measurements from spectroscopy and photometry – blue camera observations of HD 169142 show differences up to 50%. The cause of the difference is currently unknown, and may be due to the large systematic flux uncertainties. We will revisit this issue with future revisions of the instrumental calibration and reduction techniques. As the 70 μ m flux from the photometric observation is more consistent with photometry in the literature, we consider it the more reliable of the two flux values, and note the [OI] line flux for HD 169142 may be higher by up to 50%.

3. Discussion

3.1. Discs or outflows?

A primary driver of the GASPS project is to use far-infrared emission lines to characterize the gas content of circumstellar discs. However, many young star with accreting discs are also known to drive outflows. Their shocks produce strong line emission in this band (Liseau et al. 2006). Following Hollenbach (1985), we can bound the mass outflow rate from the [OI] 63 luminosity, $\dot{M}_{\text{outflow}}(M_{\odot}/\text{yr}) \approx 10^{-4} L_{[\text{OI}]63}(L_{\odot})$. As some shock emission may emerge through other lines, for the observed fluxes, this translates into lower limits of 46, 3.5, and $9.1 \times 10^{-10} M_{\odot}/\text{yr}$ for HD 169142, TW Hya, and RECX 15 respectively. Comparison to observed accretion rates (Sect. 1) rule out an outflow shock origin for the emission from HD 169142, and requires the outflow rate to be a very high fraction (≥ 0.2) of the accretion rate for TW Hya and RECX 15. This is much greater than the ratio found in T Tauri stars by Hartigan et al. (1995).

Further, there is no direct evidence for outflows in these sources. We do not find [OI] in other IFU spaxels in excess of that expected in the wings of the point-spread function, which places an upper bound to the extent of the emission to less than 680, 260, and 460 AU for HD169142, TW Hya, and RECX 15 respectively. With the previously noted exception of RECX 15, the spectra are not resolved, precluding jets with line-of-sight velocities greater than $\sim 45 \text{ km s}^{-1}$. As noted in Sect. 2.1, the broad profile in RECX 15 is suspected to be due to instrumental effects. This also agrees with the narrow spectral profiles of OI 6300 Å and CO rotational lines in HD169142 and TW Hydra (Acke et al. 2005; Raman et al. 2006; Qi et al. 2004).

While J and C shock models predict $\frac{[\text{OI}]63}{[\text{OI}]145}$ and $\frac{[\text{OI}]63}{[\text{CII}]158}$ line ratios ≥ 10 (e.g. Hollenbach & McKee 1989), the analysis of 500 archival ISO observations by Liseau et al. (2006) showed typical ratios of ~ 1 – 20 and 0.1 – 5 for background subtracted sources (locus of detections of Liseau et al. from their Fig. 2, are shown in Fig. 2, along with models from the DENT grid as discussed below). The high $\frac{[\text{OI}]63}{[\text{CII}]158}$ ratios of our targets (≥ 10 for HD 169142, and ≥ 5 for TW Hya and RECX 15) only marginally overlap the outflow locus – less than 10% of the Liseau et al. sources have line ratios large enough to be consistent with our detections.

We compare our results with a subset of the DENT grid of models, selected to match the observed range of [OI] 63 fluxes (scaled to 140 pc). This includes all models with [OI] 63 line strengths from $(4\text{--}90) \times 10^{-18} \text{ W/m}^2$. We show the locus of these models in Fig. 2, which include models with disc masses ranging from 10^{-6} to $10^{-1} M_{\odot}$. Location within this plot can generally constrain disc gas mass to within an order of magnitude; further constraints using detected line strengths, stellar properties and data from other sources such as optical and mm continuum and line measurements, can refine this estimate and allow for determination of the properties of the originating structure (see Kamp et al. 2010; Woitke et al. 2010). Predicted line flux ratios of the best fit models for HD 169142 (Meeus et al. 2010), TW Hya (Thi et al. 2010), and for RECX 15 (Woitke et al., in prep) indicate that with somewhat longer observations, TW Hya and RECX 15 may be detectable in [CII] 158, whereas HD 169142 may be detectable in [OI] 145.

Given the lack of high velocity spectral observations, poor overlap with observed line ratios of jet sources, and consistency of line ratios with disc models, we cautiously interpret our results in the context of disc emission.

3.2. [CII] non-detections

The lack of detections of the [CII] 158 line from any of our SDP targets is surprising. Kamp et al. (2010) predict [CII] 158 emission of $\sim 10^{-17} \text{ W/m}^2$ for HAe discs with radii $\sim 200 \text{ AU}$,

³ PACS_ScanMap_ReleaseNote_23Feb2010.pdf

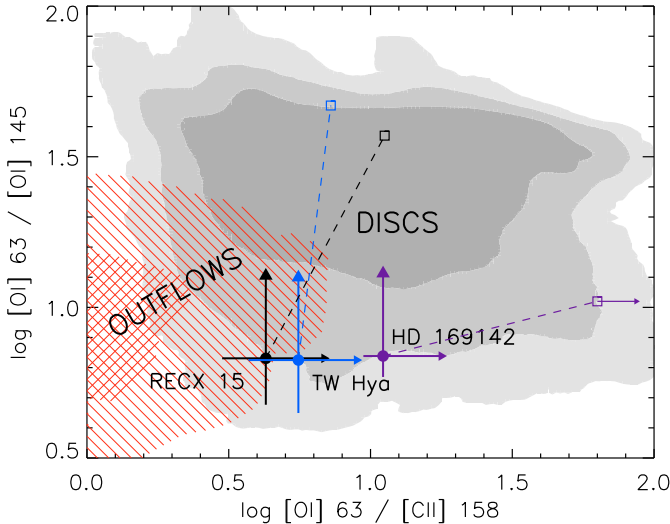


Fig. 2. Line flux ratio diagram comparing models with observations. Dots represent GASPS observations with lines indicating 3σ limits, and arrows indicating lower limits. Squares show predicted line flux ratios based on the detailed models (HD 169142: Meeus et al. 2010; TW Hya: Thi et al. 2010; RECX 15: Woitke et al. (in prep.). Outflow sources are shown in red (Liseau et al. 2006, lines indicate contour containing 90% of points, hash indicates 68% of points). Gray regions indicate 99, 90, and 68% of DENT models with [OI] 63 flux consistent with our observations ($4\text{--}90 \times 10^{-18} \text{ W/m}^2$ at 140 pc).

with emission varying with disc surface area. HD 169142 is known from mm imaging to have a disc of radius >200 AU (Raman et al. 2006). Other factors must be at play, such as flat discs, lack of UV flux, or shielding by a large amount of small dust grains at large radii.

3.3. A new regime of disc studies

The ratio of [OI] 63 line strength to continuum varies substantially, from 0.3–1 for HD 169142 and TW Hya to 17 for RECX 15. Both line and continuum are likely optically thick and the stellar luminosities are very different for the three sources but one possible explanation for the strikingly high line-to-continuum ratio in RECX 15 is an enhanced gas-to-dust ratio (see Fig. 3 of Pinte et al. 2010).

Further, whereas HD 169142 and TW Hya have strong mm continuum fluxes indicating large dust masses, RECX 15 is undetected at $850 \mu\text{m}$ (Phillips, in prep.). This is similar to the 49 Ceti disc which is detected at 1.3 mm in CO 2–1 but not in the continuum (Hughes et al. 2008). These comparisons suggest a high gas-to-dust ratio; however, further modeling is necessary (see Woitke et al., in prep.). The GASPS survey will reveal how common are discs with gas emission yet lacking detectable dust.

Prior studies of circumstellar disc gas have focused on near-infrared emission lines from the hot inner disc (e.g. Brittain et al. 2007) or mm lines from the cool outer disc (e.g. Dent et al. 2005). Far-infrared line emission from intermediate radii of gas rich discs has not yet been explored in similar detail. Thanks to *Herschel*'s unprecedented sensitivity, GASPS will examine the gas content of circumstellar discs in previously inaccessible regimes.

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References

- Acke, B., van den Ancker, M. E., & Dullemond, C. P. 2005, *A&A*, 436, 209
 Barrado Y Navascués, D. 2006, *A&A*, 459, 511
 Brittain, S. D., Simon, T., Najita, J. R., & Rettig, T. W. 2007, *ApJ*, 659, 685
 de la Reza, R., Torres, C., Quast, G., Castilho, B., & Vieira, G. 1989, *ApJ*, 343, L61
 Dent, W. R. F., Greaves, J. S., & Coulson, I. M. 2005, *MNRAS*, 359, 663
 Dunkin, S. K., Barlow, M. J., & Ryan, S. G. 1997, *MNRAS*, 286, 604
 Gautier, III, T. N., Rebull, L. M., Stapelfeldt, K. R., & Mainzer, A. 2008, *ApJ*, 683, 813
 Grady, C. A., Schneider, G., & Hamaguchi, K., et al. 2007, *ApJ*, 665, 1391
 Hartigan, P., Edwards, S., & Ghandour, L. 1995, *ApJ*, 452, 736
 Herczeg, G. J., Wood, B. E., Linsky, J. L., Valenti, J. A., & Johns-Krull, C. M. 2004, *ApJ*, 607, 369
 Hollenbach, D. 1985, *Icarus*, 61, 36
 Hollenbach, D., & McKee, C. F. 1989, *ApJ*, 342, 306
 Hughes, A. M., Wilner, D. J., Calvet, N., et al. 2007, *ApJ*, 664, 536
 Hughes, A. M., Wilner, D. J., Kamp, I., & Hogerheijde, M. R. 2008, *ApJ*, 681, 626
 Kamp, I., Tilling, I., Woitke, P., Thi, W.-F., & Hogerheijde, M. 2010, *A&A*, 510, 18
 Kaufman, M. J., Wolfire, M. G., Hollenbach, D. J., & Luhman, M. L. 1999, *ApJ*, 527, 795
 Lawson, W. A., Crause, L. A., Mamajek, E. E., & Feigelson, E. D. 2002, *MNRAS*, 329, L29
 Lawson, W. A., Lyo, A.-R., & Muzerolle, J. 2004, *MNRAS*, 351, L39
 Liseau, R., Justtanont, K., & Tielens, A. G. G. M. 2006, *A&A*, 446, 561
 Lorenzetti, D., Tommasi, E., Giannini, T., et al. 1999, *A&A*, 346, 604
 Mamajek, E. E., Lawson, W. A., & Feigelson, E. D. 2000, *ApJ*, 544, 356
 Meeus, G., Pinte, C., Woitke, P., et al. 2010, *A&A*, 518, L124
 Muzerolle, J., Calvet, N., Briceño, C., Hartmann, L., & Hillenbrand, L. 2000, *ApJ*, 535, L47
 Nordström, B., Mayor, M., Andersen, J., et al. 2004, *A&A*, 418, 989
 Ott, S. 2010, in *Astronomical Data Analysis Software and Systems XIX*, ed. Y. Mizumoto, K.-I. Morita, & M. Ohishi, ASP Conf. Ser., in press
 Panić, O., Hogerheijde, M. R., Wilner, D., & Qi, C. 2008, *A&A*, 491, 219
 Pascucci, I., Hollenbach, D., Najita, J., et al. 2007, *ApJ*, 663, 383
 Perrymann, M. A. C., Lindegren, L., Kovalevsky, J., et al. 1997, *A&A*, 323, L49
 Pilbratt, G. L., et al. 2010, *A&A*, 518, L1
 Pinte, C., Ménard, F., Duchêne, G., & Bastien, P. 2006, *A&A*, 459, 797
 Pinte, C., et al. 2010, *A&A*, 518, L126
 Poglitsch, A., et al. 2010, *A&A*, 518, L2
 Qi, C., Ho, P. T. P., Wilner, D. J., et al. 2004, *ApJ*, 616, L11
 Raman, A., Lisanti, M., Wilner, D. J., Qi, C., & Hogerheijde, M. 2006, *AJ*, 131, 2290
 Thi, W.-F., Mathews, G. S., Ménard, F., et al. 2010, *A&A*, 518, L125
 van Boekel, R., Min, M., Waters, L. B. F. M., et al. 2005, *A&A*, 437, 189
 Wichmann, R., Bastian, U., Krautter, J., Jankovics, I., & Rucinski, S. M. 1998, *MNRAS*, 301, L39
 Woitke, P., Kamp, I., & Thi, W.-F. 2009, *A&A*, 501, 383
 Woitke, P., Pinte, C., Tilling, I., et al. 2010, *MNRAS*, 405, L26
 Zuckerman, B., & Song, I. 2004, *ApJ*, 603, 738

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